Symmetry Energy Parameters and How the EOS is Taking Shape

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July 2, 2012

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INT Supernova and Nucleosynthesis Workshop
Seattle 25 June – 27 July, 2012

Outline

- Why the Symmetry Energy is Important for Neutron Stars
- Constraints on Symmetry Parameters from Nuclear Experiments
  - Binding Energies
  - Heavy ion Collisions
  - Neutron Skin Thicknesses
  - Dipole Polarizabilities
  - Giant (and Pygmy) Dipole Resonances
- Theoretical Calculations of Pure Neutron Matter and Their Predictions for the Symmetry Energy
- Astrophysical Observations
  - Mass Measurements of Neutron Stars in Binaries
  - Simultaneous Mass and Radius Measurements
    - Thermal Emission from Cooling Neutron Stars
    - Photospheric Radius Expansion X-Ray Bursters
    - Pulse Shape Modeling of Accreting Millisecond Pulsars
  - The Inferred Universal Mass-Radius Relation and the Neutron Star EOS

Mass-Radius Diagram and Theoretical Constraints

GR:

$R > 2GM/c^2$

$P < \infty$:

$R > (9/4)GM/c^2$

$M < M_{max}$

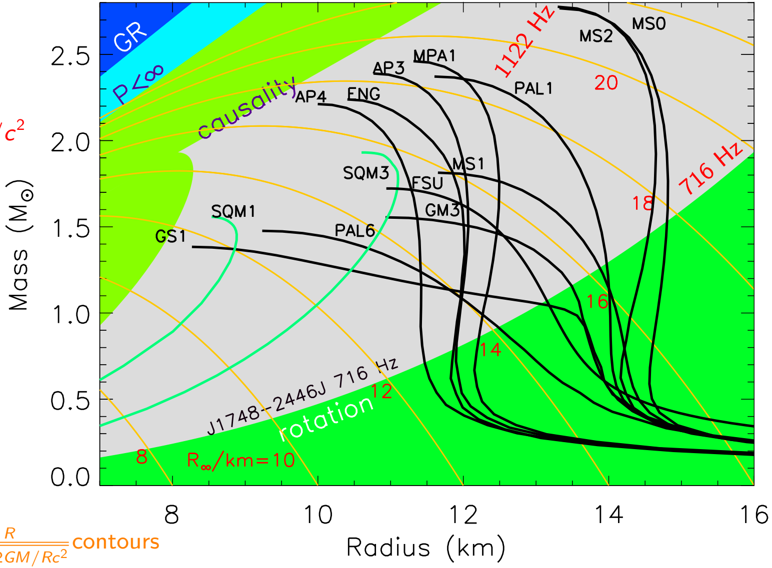
causality:

$R \gtrsim 2.9GM/c^2$

— normal NS

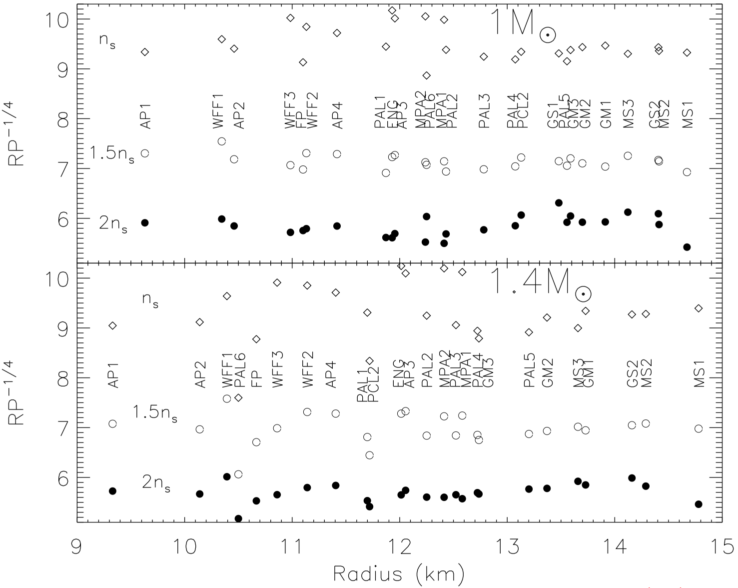
— SQS

— $R_\infty = \frac{R}{\sqrt{1-2GM/Rc^2}}$ contours



The Radius – Pressure Correlation

$R \propto p^{1/4}$



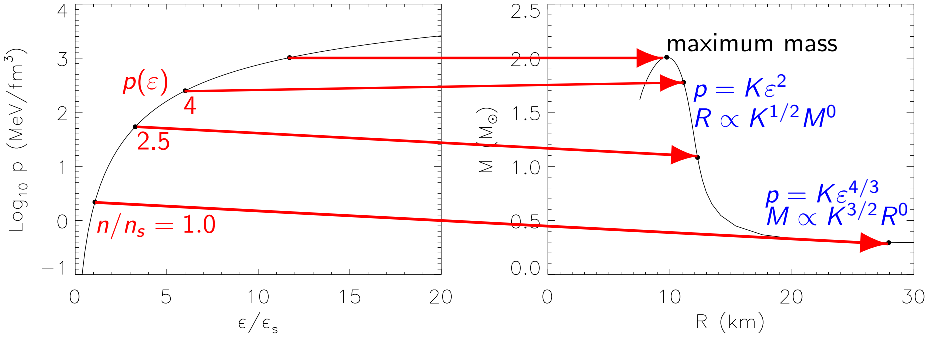
Lattimer & Prakash (2001)

Neutron Star Structure

Tolman-Oppenheimer-Volkov equations

$$\frac{dp}{dr} = -\frac{G}{c^4}\frac{(mc^2 + 4\pi pr^3)(\varepsilon + p)}{r(r - 2Gm/c^2)}; \quad \frac{dm}{dr} = 4\pi\frac{\varepsilon}{c^2}r^2$$

Newtonian Polytropes: $p = K\varepsilon^\gamma$; $\quad M \propto K^{1/(2-\gamma)}R^{(4-3\gamma)/(2-\gamma)}$



Nuclear Symmetry Energy

Defined as the difference between energies of pure neutron matter ($x = 0$) and symmetric ($x = 1/2$) nuclear matter

$$S(\rho) = E(\rho, x = 0) - E(\rho, x = 1/2)$$

Expanding around saturation density (ρ_s) and symmetric matter

$$E(\rho, x) = E(\rho, x = 1/2) + (1 - 2x)^2 E_{sym}(\rho) + \ldots$$
$$E_{sym}(\rho) = S_v + \frac{L}{3}\frac{\rho - \rho_s}{\rho_s} + \frac{K_{sym}}{18}\left(\frac{\rho - \rho_s}{\rho_s}\right)^2 + \ldots$$
$$S_v = \frac{1}{8}\frac{\partial^2 E}{\partial x^2}\bigg|_{\rho_s,1/2}, \quad L = \frac{3}{8}\frac{\partial^3 E}{\partial\rho\partial x^2}\bigg|_{\rho_s,1/2}, \quad K_{sym} = \frac{9}{8}\frac{\partial^4 E}{\partial\rho^2\partial x^2}\bigg|_{\rho_x,1/2}$$

Thus, $E_{sym}(\rho) \simeq S(\rho)$ if higher-than-quadratic terms are small. Can be connected to neutron matter:

$$S(\rho_s) = E(\rho_s, x = 0) + B \approx S_v, \qquad p(\rho_s, x = 0) \approx L\rho_s/3$$

$R \propto p(\rho_s - 2\rho_s, x = 0)^{1/4}$ **(Lattimer & Prakash 2001)**

Nuclear Binding Energies

$$E_{sym}(N,Z) = I^2(S_v A - S_s A^{2/3})$$

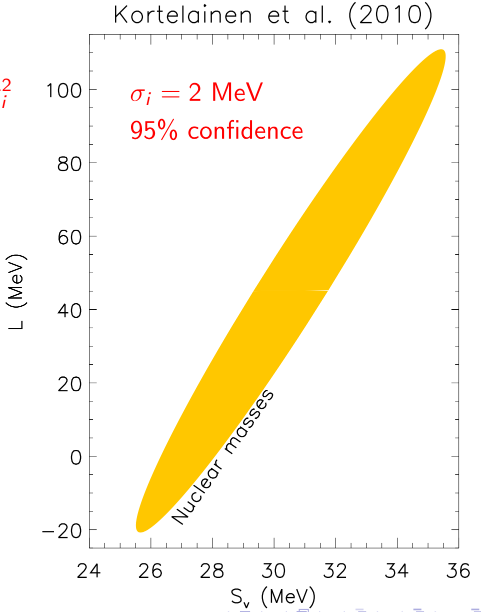
$$\chi^2 = \mathcal{N}^{-1}\sum_i (E_{ex,i} - E_{sym,i})^2/\sigma_i^2$$
$$\chi_{vv} = \frac{2}{\mathcal{N}}\sum_i I_i^4 A_i^2$$
$$\chi_{ss} = \frac{2}{\mathcal{N}}\sum_i I_i^4 A_i^{4/3}$$
$$\chi_{vs} = \frac{2}{\mathcal{N}}\sum_i I_i^4 A_i^{5/3}$$

$$\sigma_{S_v} = \sqrt{\frac{2\chi_{ss}}{\chi_{vv}\chi_{ss}-\chi_{sv}^2}}$$
$$\sigma_{S_s} = \sqrt{\frac{2\chi_{vv}}{\chi_{vv}\chi_{ss}-\chi_{sv}^2}}$$

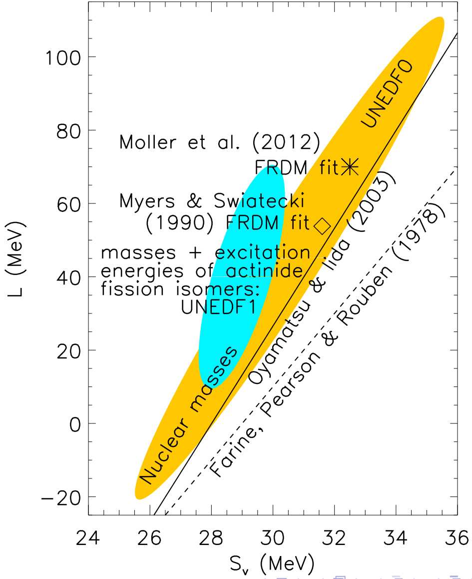
$$\alpha = \frac{1}{2}\tan^{-1}\frac{2\chi_{vs}}{\chi_{vv}-\chi_{ss}}$$
$$r_{vs} = -\frac{\chi_{vs}}{\sqrt{\chi_{vv}\chi_{ss}}}$$

$$S_s \approx 0.95 S_v + 0.65 L$$

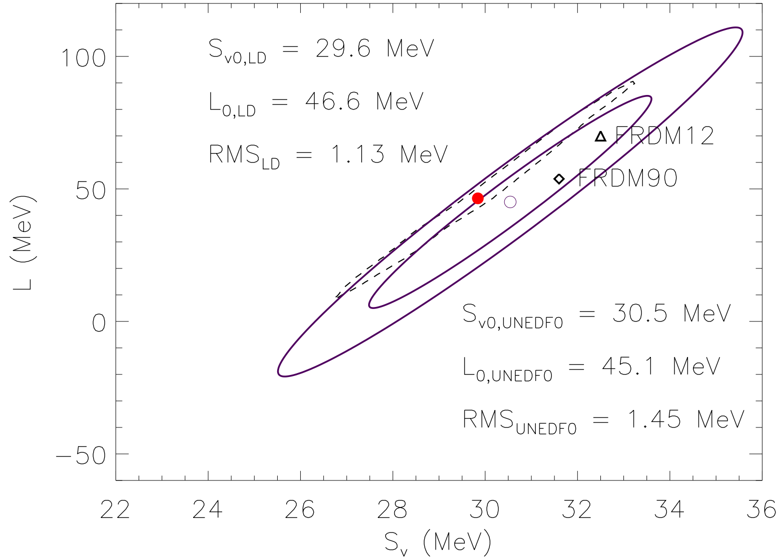
$$E_{sym}(N,Z) = \frac{S_v A I^2}{1+(S_s/S_v)A^{-1/3}}$$



Nuclear Binding Energies



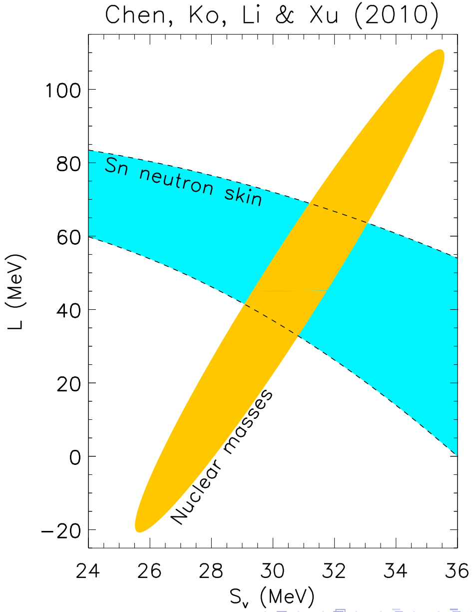
Comparison of Microscopoic and Liquid Droplet Mass Fits



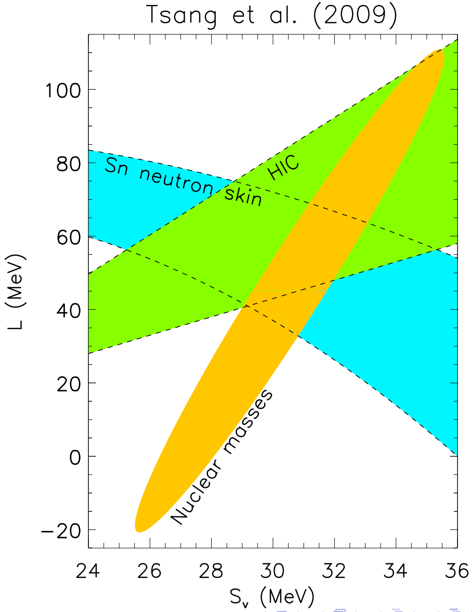
Neutron Skin Thickness

$$R_n - R_p \simeq \sqrt{3/5}\, t_{np}$$

$$t_{np} = \frac{2r_o}{3} \frac{S_s I}{S_v + S_s A^{-1/3}}$$



Heavy Ion Collisions

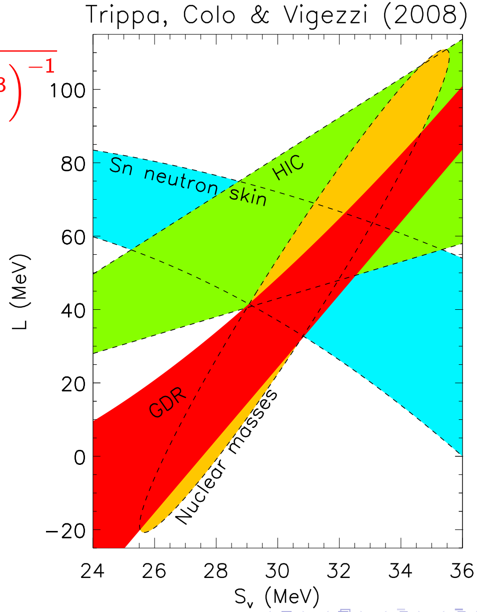


Giant Dipole Resonances

$$E_{-1} \propto \sqrt{S_v \left(1 + \frac{5S_s}{3S_v} A^{-1/3}\right)^{-1}}$$

Correlation between E_{-1} and E_{sym} maximized when

$$E_{sym,208}/AI^2 = \frac{S_v}{1+(S_s/S_v)A^{-1/3}} \simeq S(\rho = 0.1)$$

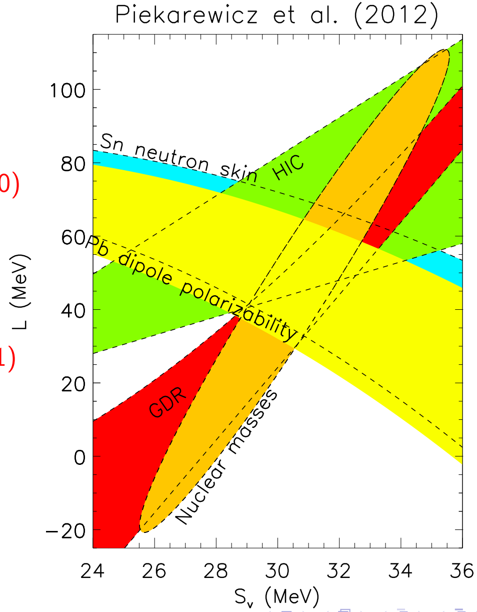


Dipole Polarizability

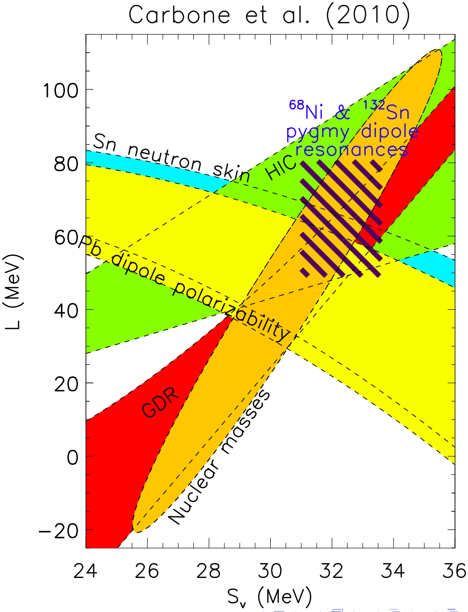
α_D and $R_n - R_p$ in ^{208}Pb are 98% correlated

Reinhard & Nazawericz (2010)

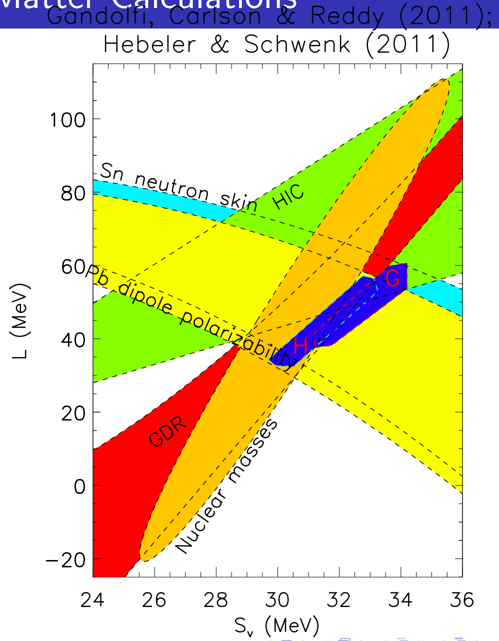
Data from Tamii et al. (2011)



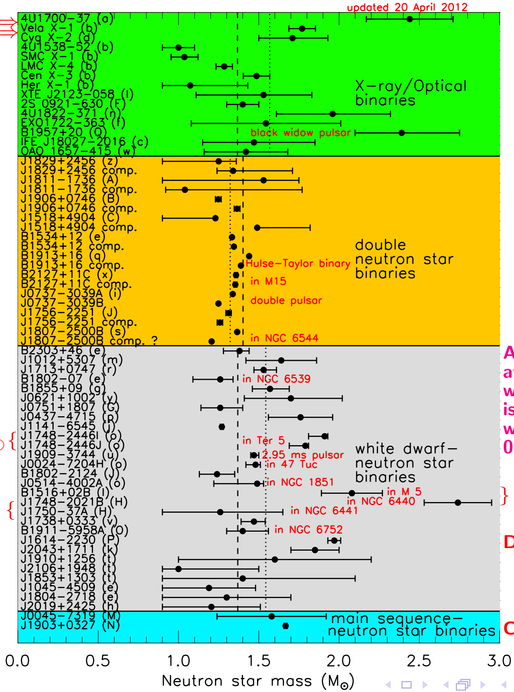
Pygmy Dipole Resonances



Theoretical Neutron Matter Calculations



Black hole? ⇛
Firm lower mass limit? ⇛



Although simple average mass of w.d. companions is 0.27 M☉ larger, weighted average is 0.08 M☉ smaller

$M > 1.68$ M☉ { **95% confidence**

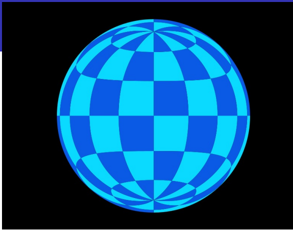
} **w.d. companion? statistics?**

Freire et al. 2007 {

Demorest et al. 2010

Champion et al. 2008

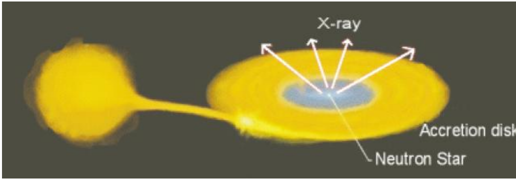
Simultaneous Mass/Radius Measurements



- The measurement of flux and temperature yields an apparent angular size (pseudo-BB):

$$\frac{R_\infty}{D} = \frac{R}{D} \frac{1}{\sqrt{1 - 2GM/Rc^2}}$$

- Observational uncertainties include distance, interstellar absorption (UV and X-rays), atmospheric composition

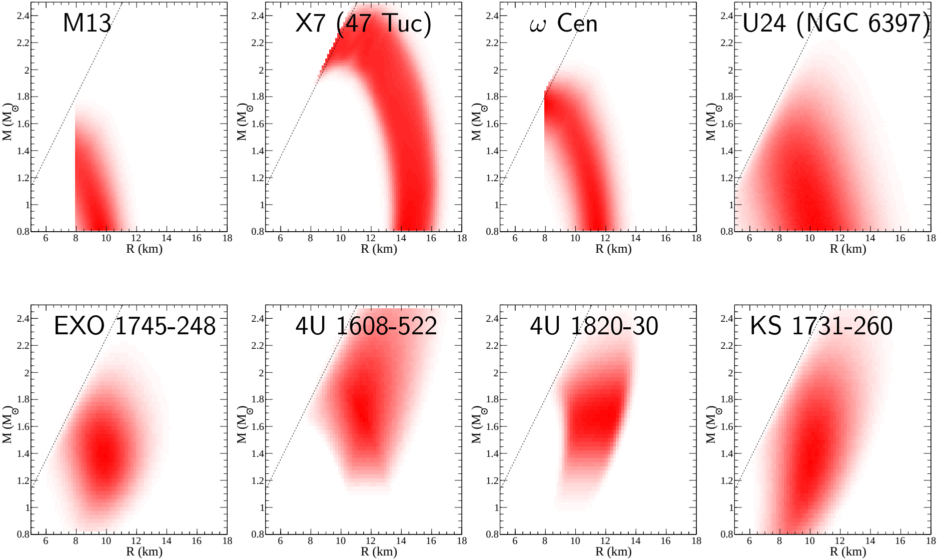


Best chances for accurate radius measurement:

- Nearby isolated neutron stars with parallax (uncertain atmosphere)
- Quiescent X-ray binaries in globular clusters (reliable distances, low B H-atmospheres)
- Bursting sources with peak fluxes close to Eddington limit (where gravity balances radiation pressure)

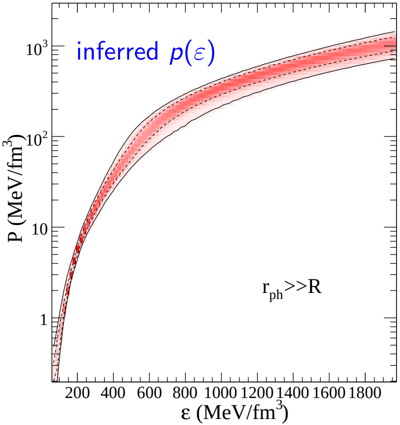
$$F_{Edd} = \frac{cGM}{\kappa D^2}$$

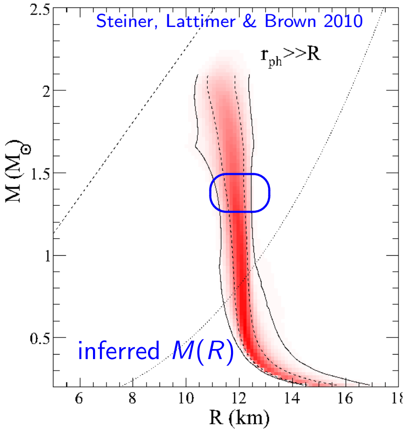
$M - R$ Probability Estimates



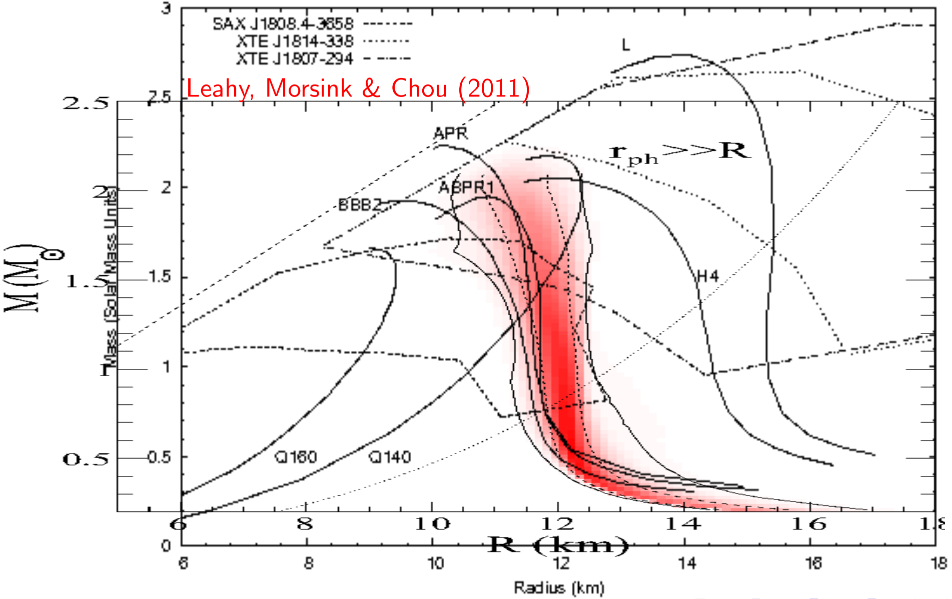
Bayesian TOV Inversion

- $\varepsilon < 0.5\varepsilon_0$: Known crustal EOS
- $0.5\varepsilon_0 < \varepsilon < \varepsilon_1$: EOS parametrized by K, K', S_v, γ
- Polytropic EOS: $\varepsilon_1 < \varepsilon < \varepsilon_2$: n_1; $\varepsilon > \varepsilon_2$: n_2
- EOS parameters $K, K', S_v, \gamma, \varepsilon_1, n_1, \varepsilon_2, n_2$ uniformly distributed
- $M_{max} \geq 1.97$ M$_\odot$, causality enforced
- All sources equally weighted

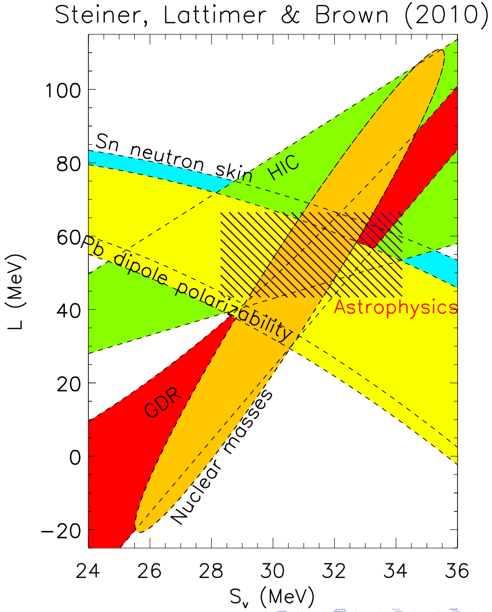




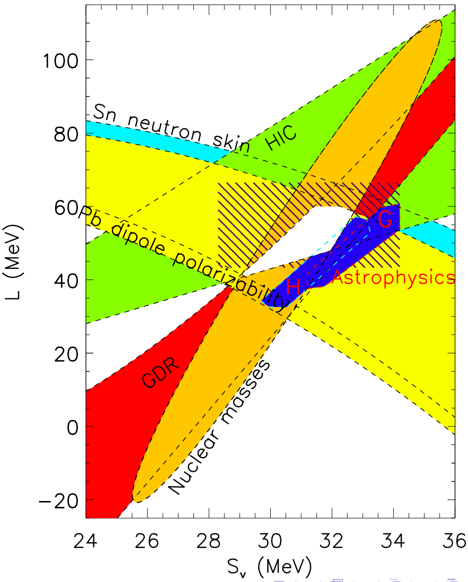
Pulse Shape Modeling



Astronomical Observations



Combined Constraints



Consistency with Neutron Matter and Heavy-Ion Collisions